

Discovery of the Most Ultra-Luminous QSO Using *GAIA*, SkyMapper, and WISE

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Abstract

We report the discovery of the ultra-luminous quasi-stellar object SMSS J215728.21–360215.1 with magnitude $z = 16.9$ and $W4 = 7.42$ at redshift 4.75. Given absolute magnitudes of $M_{145,AB} = -29.3$, $M_{300,AB} = -30.12$, and $\log L_{bol}/L_{bol, \odot} = 14.84$, it is the quasi-stellar object with the highest unlensed UV-optical luminosity currently known in the Universe. It was found by combining proper-motion data from *Gaia* DR2 with photometry from SkyMapper DR1 and the Wide-field Infrared Survey Explorer. In the *GAIA* database, it is an isolated single source and thus unlikely to be strongly gravitationally lensed. It is also unlikely to be a beamed source as it is not discovered in the radio domain by either NRAO-VLA Sky Survey or Sydney University Molonglo Southern Survey. It is classed as a weak-emission-line quasi-stellar object and possesses broad absorption line features. A lightcurve from ATLAS spanning the time from 2015 October to 2017 December shows little sign of variability.

Keywords: galaxies: active – galaxies: high-redshift – quasars: general

1 INTRODUCTION

Black holes at the centres of galaxies reach masses of over 10 billion times that of our Sun. Surprisingly, we have found such massive black holes already in the early Universe, just 800 million years after the Big Bang (Wu et al. 2015). How they grew to such mass so early after the Big Bang is a profound puzzle for physics. They must have grown at super-Eddington rates for a long period of time, or they originate from massive seed black holes that formed during the dark early ages by direct collapse (Bromm & Loeb 2003; Pacucci, Volonteri, & Ferrara 2015).

Currently, we can only discover such super-massive black holes in the distant early universe while they grow rapidly and accrete vast amounts of matter. This makes them appear as very luminous quasi-stellar objects (QSOs, Schmidt 1963) when we have a clear view of the accretion disk around the black hole, and as type-2 QSOs and infrared-luminous galaxies when that view is blocked by dust (e.g. Seymour et al. 2007; Lacy et al. 2013); at present, we can only measure the masses of their black holes when we have a clear view.

Finding the most luminous, optically bright, QSOs is important for several reasons: (i) they point us to the most massive black holes that pose the greatest challenge to any physical growth scenario; (ii) they ionise large volumes of neutral gas around them and contribute to cosmic reionisation (Fan, Carilli, & Keating 2006a; Wu et al. 2015); (iii) they reveal the metal enrichment in the early universe by shining like beacons through the gas content of high-redshift galaxies along the line-of-sight that are otherwise hard to observe (Ryan-Weber et al. 2009; Simcoe et al. 2011); and (iv) they will eventually enable the most sensitive direct observations of the expansion of the Universe (Liske et al. 2008).

Among optically luminous QSOs at high redshift, the most impressive objects are J0306+1853 at $z = 5.36$, which is powered by a black hole of 10 billion solar masses (Wang et al. 2015), and J0100+2802 at $z = 6.3$ with a black hole of 12 billion M_{\odot} (Wu et al. 2015); both of them radiate at their Eddington limit. Tsai et al. (2015) claim the bolometrically most luminous object in the Universe is the ELIRG

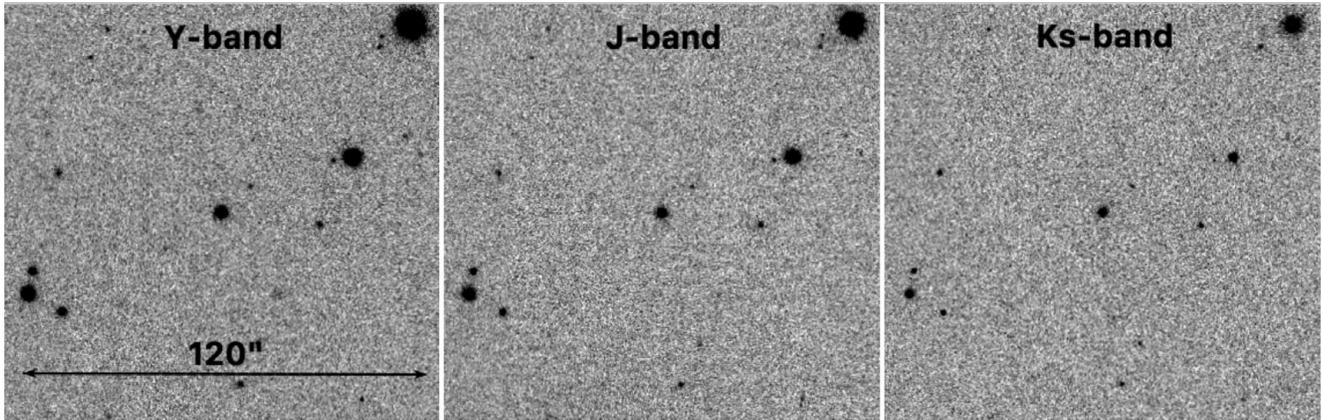


Figure 1. Images of J2157–3602 from the VISTA Hemispheric Survey (VHS) in YJK_s -bands. The QSO appears as an isolated, single point source just like it does in the *GAIA* R_p -band, so its brightness is unlikely to be boosted by strong gravitational lensing.

WISE J224607.57-052635.0 at $z = 4.593$ with $\log L_{\text{bol}}/L_{\odot} = 14.54$ but disagree with Wu et al. (2015) on the bolometric correction.

Here, we report a newly identified high-redshift QSO with the highest unlensed UV-optical luminosity known at present. Throughout the paper, 2MASS and WISE magnitudes are used in the Vega system, while optical magnitudes are in the AB system (Oke & Gunn 1983). We adopt a standard flat Λ CDM cosmology with parameters $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$, and $\Omega_{\Lambda} = 0.7$.

2 Data

2.1. Selection and photometry

The discovery of SMSS J215728.21–360215.1 (hereafter, J2157–3602) was helped by the Data Release 2 from *GAIA* (Gaia Collaboration et al. 2018) and its unprecedented precision on proper motion measurements. Prior to *GAIA* DR2, ultra-luminous QSOs were discovered only serendipitously, because QSO candidate lists were inflated by staggering contamination from cool stars. Even though QSO colours differ from star colours to some degree, the scatter in star colours leads to overlapping loci in colour space. And at bright levels, where the luminosity function of high- z QSOs plummets towards zero space density, stars vastly outnumber real QSOs. Thus, spectroscopic follow-up was often biased against the brightest and potentially most exciting candidates as they were by far the least likely to be QSOs.

However, the main contaminants of the search box for high- z QSOs are nearby Galactic stars of low mass and temperature, which can now be identified very reliably from the proper motions in the *GAIA* database. Hence, immediately after the release of *GAIA* DR2 on 2018 April 25, we searched for red objects using $B_p - R_p$ in the *GAIA* database that are consistent with having no significant proper motion. Candidates were then cross-matched against SkyMapper DR1

(Wolf et al. 2018), the 2 Micron All Sky Survey (2MASS, Skrutskie et al. 2006) and the Wide-field Infrared Survey Explorer (WISE, Wright et al. 2010), and further selected to have optical-infrared colours appropriate for high- z QSOs (details will be published by Wolf et al., in preparation).

J2157–3602 appears in *GAIA* DR2 as a single isolated source; the nearest neighbour detected by *GAIA* is ~ 42 arcsec away, while the resolution achieved by *GAIA* reaches up to ~ 0.1 arcsec¹. *GAIA* is complete to magnitude $R_p \approx 19$, while our object is ~ 2 mag brighter; this suggests that the brightness of the object is not boosted by strong gravitational lensing. We find the same conclusion after searching infrared images from the VISTA Hemisphere Survey (VHS, McMahon et al. 2013) for signs of an extended source that could signify either multiple images or the presence of a lensing galaxy. In all three bands, YJK_s , the source is fully consistent with a point source, where the PSF FWHM is 1.04, 0.94, and 0.91 arcsec, respectively (see also Figure 1). We checked the Abell, Corwin, & Olowin (1989) catalogue of galaxy clusters in a 2° radius around the QSO and found no entry that could have weakly lensed it.

J2157–3602 is clearly detected by SkyMapper in the far-red bands with $i = 17.37$ and $z = 17.11$, but not in any of the bluer bands, which are only complete to 18 mag in DR1. It was detected by VST-ATLAS (Shanks et al. 2015) with $g \approx 21$, $r = 18.68$, $i = 17.32$, and $z = 16.93$. Most photometry for high- z QSOs is from the SDSS survey (see, e.g., Wang et al. 2016), and while VST filters are very similar to the SDSS, the SkyMapper passbands are different. The optical colour is consistent with the $4.7 < z < 5.4$ QSO selection criteria used by Wang et al. (2016), but it is outside of the optical selection box used by Richards et al. (2002).

It was also clearly detected in 2MASS and WISE, reaching magnitude 7.42 in the $22 \mu\text{-band}$ W4. The WISE position is consistent with the *GAIA* position within 0.05 arcsec. The

¹ In Qi et al. (2015) the object was listed with an incompatible, large proper motion of $\mu_{\text{dec}} = -11.3 + / - 1.5$ milliarcsec.

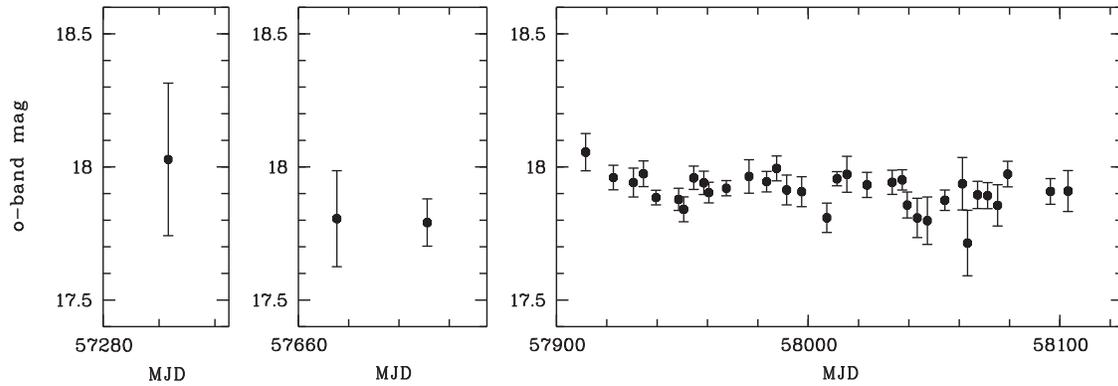


Figure 2. Light curve of J2157–3602 from the Asteroid Terrestrial-impact Last Alert System (ATLAS) project covering three seasons from 2015 October to 2017 December, after binning magnitudes into 5-d intervals to reduce noise. There is no sign of strong variability, but slow variations at a ± 0.1 mag level can be seen.

nearest neighbouring source in the WISE catalogue is 26 arcsec away and is 4 mag fainter in W1 and 1 mag fainter in W4. With a colour of $W1-W2 = 0.51 \pm 0.03$, the QSO was at the edge of the selection box used by Wang et al. (2016). Their selection of $W1-W2 > 0.5$ was motivated by suppressing contamination from stars, but with *GAIA*, it is feasible to push the selection limits further into the stellar locus.

In Figure 2, we show the light curve of the object in the orange filter of the Asteroid Terrestrial-impact Last Alert System (ATLAS, Tonry et al. 2018) covering three observing seasons from 2015 October to 2017 December. The mean error and scatter is ~ 0.1 mag and little variability is detected above that, but some variation can be seen after binning the data into 5-d intervals.

J2157–3602 is not detected in any large radio survey: neither by the NRAO-VLA Sky Survey (NVSS, Condon et al. 1998), which has a flux limit of 2.5 mJy at 1.4 GHz, nor by the Sydney University Molonglo Southern Survey (SUMSS, Bock, Large, & Sadler 1999) with a limit of 5 mJy at 843 MHz. We use the radio-to-optical ratio $R = f_{v,6\text{cm}}/f_{v,B}$ as defined by Kellermann et al. (1989) to label the object as radio-loud ($R > 10$) or radio-quiet ($R < 1$); in that work, the ratio is based on the flux density at 6 cm and 440 nm wavelength for a sample with an average redshift of ~ 0.5 , but for our object, these bands are redshifted to ~ 23 cm, which is approximately the NVSS wavelength, and 1710 nm. We find a ratio of $R < 2.0$ and thus assume that the brightness is not significantly increased by the presence of a jet with relativistic boosting. Photometric and additional properties of the object are listed in Table 1.

2.2. Spectroscopy

We took a spectrum of this object during the full-moon night of 2018 April 29 18:09 UT, using the ANU 2.3-m telescope at Siding Spring Observatory in Australia with the Wide Field Spectrograph (WiFeS, Dopita et al. 2010). We took two exposures of 600 s with the WiFeS R3000 grating that covers

Table 1. Properties of QSO SMSS J215728.21–360215.1.

Gaia positional data	
RA ($^{\circ}$)	329.3676229
Dec ($^{\circ}$)	−36.0375573
b_{gal}	−52.2195997
l_{gal}	8.2221573
parallax (milli-arcsec)	0.15 ± 0.33
proper motion (milli-arcsec yr $^{-1}$)	0.03 ± 0.56
ref-epoch	2015.5
$E(B - V)_{\text{SFD}}$	0.0147
Photometry	
Gaia B_p	19.794 ± 0.058
Gaia G	18.286 ± 0.003
Gaia R_p	17.136 ± 0.007
VST-ATLAS μ_{AB}	undetected
VST-ATLAS g_{AB}	20.94 ± 0.03
VST-ATLAS r_{AB}	18.68 ± 0.01
VST-ATLAS i_{AB}	17.32 ± 0.01
VST-ATLAS z_{AB}	16.93 ± 0.01
SkyMapper r_{AB}	undetected
SkyMapper i_{AB}	17.37 ± 0.02
SkyMapper z_{AB}	17.11 ± 0.02
VHS Y	16.15 ± 0.01
VHS J	15.65 ± 0.01
VHS K_s	14.25 ± 0.01
2MASS J	15.70 ± 0.06
2MASS H	14.84 ± 0.05
2MASS K	14.16 ± 0.06
WISE W1 (mpro)	13.11 ± 0.02
WISE W2 (mpro)	12.60 ± 0.03
WISE W3 (mpro)	9.91 ± 0.05
WISE W4 (mpro)	7.42 ± 0.13
NVSS $f_{1.4\text{GHz}}$	< 2.5 mJy
SUMSS $f_{843\text{MHz}}$	< 5.0 mJy
Redshift and luminosities	
Redshift	~ 4.75
$t_{\text{look-back}}$	12.23 Gyr
$M_{145,\text{AB}}$	−29.30
$M_{300,\text{AB}}$	−30.12
M_{bol}	−32.36
L_{bol}/L_{\odot}	6.95×10^{14}

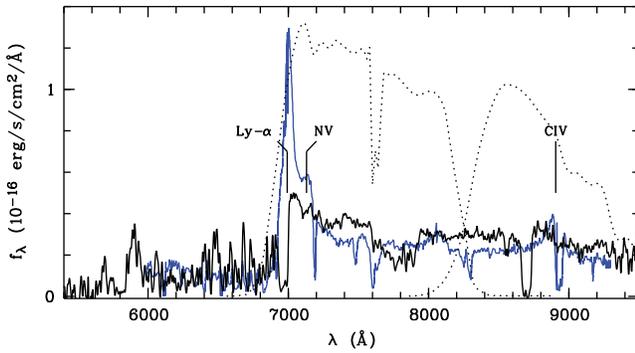


Figure 3. Spectrum of J2157–3602 (thick line) compared to PMN J1451 (thin line). Strong Ly- α emission makes PMN J1451 brighter in VST i -band (dashed), while it is fainter in the continuum and z -band. Both spectra are calibrated from photometry.

the wavelength range from 5 300 to 9 800 Å with a resolution of $R = 3000$. The data were reduced using the Python-based pipeline PyWiFeS (Childress et al. 2014). The flux density was calibrated using the standard stars HD 55496 and HD 111786 observed on the same night.

The spectrum is shown in Figure 3 and shows only weak emission lines but a clear Lyman- α forest. The object is thus classified as a Weak-emission-line QSO (WLQ, see Diamond-Stanic et al. 2009). Hints of emission lines are seen for Lyman- α and NV as well as CIV, and based on these, we determine the redshift to $z = 4.75$. Using Ned Wright’s online cosmology calculator², we find a luminosity distance of $D_L = 43.91$ Gpc. Deep broad absorption lines (BALs) are observed on the blue side of the Ly- α and CIV line. For WLQs, a precise redshift determination is hampered by the poorly defined centre of the broad emission lines. An alternative redshift of $z = 4.64$ is obtained by assuming that the red edge of the CIV BAL is at the systemic velocity of the QSO.

WLQs are mostly found at high redshift, their fraction among the overall QSO population rises steeply between $z = 3$ and $z = 4$. The physical explanation for WLQs is not clear, but they are predominantly radio-quiet so their continuum is typically not boosted by radiation from relativistic jets (Diamond-Stanic et al. 2009).

3 LUMINOSITY VS. OTHER QSOS

The brightest known UV-bright QSO at $z > 4.5$ in the compilation by Wang et al. (2016) is J145147.05–151220.20 at $z = 4.763$. Previously labelled PMN J1451–1512, it was identified as a radio-loud QSO by Hook et al. (2002) with a 1.4-GHz flux density of 28.5 ± 1.0 mJy in NVSS. The object is not contained in the SDSS footprint, but in the SDSS-like filters of VST-ATLAS, it has $i = 17.22$ and $z = 17.14$. It is not listed in the SkyMapper DR1 ‘master’ table as all its data have warning flags due to proximity to a bright star (the $V \sim 7.8$ mag HD 131046, 100 arcsec away); however, visual

inspection reveals that the images are clean and we retrieve $i = 17.23 \pm 0.05$ and $z = 17.19 \pm 0.1$ from the ‘fs_photometry’ table that contains all detections.

So, PMN J1451–1512 is brighter by ~ 0.1 mag than our new object in both the i -band and in the *GAIA* R_p -band, but ~ 0.1 mag fainter in the z -band. This is related to the lack of strong emission lines and the presence of broad absorption bands in our object. PMN J1451–1512 has strong Ly- α emission with ~ 40 nm equivalent width (see Figure 3), which boosts its i -magnitude by ~ 0.3 mag, while in the continuum between the Ly- α and CIV lines, our object is 0.2 mag brighter than PMN J1451–1512.

Our object is consistently brighter than PMN J1451–1512 by 0.5 ± 0.05 mag across the whole infrared range from 1.2μ (J -band) to 22μ (W4), suggesting a nearly identical spectral slope across the rest-frame range from 200 nm to 4μ , and thus will have a larger bolometric energy output (see Figure 4). On the blue side of 200 nm rest frame, the continuum appears redder than typical for QSOs, which might have the same origin as the broad absorptions, but we leave the discussion of this point until we obtain a detailed high-S/N spectrum.

We estimate continuum luminosities at rest frame 145 and 300 nm from broadband photometry in z - and HK -bands, respectively, because our early spectrum is unfortunately too noisy and its calibration too uncertain to derive it from that; due to the weak emission and BALs, we assume this to be an appropriate estimate. We find $M_{145,AB} = -29.30$ and $M_{300,AB} = -30.12$. We translate this into monochromatic luminosities of $\nu L_{\nu, 145} = 4.71 \times 10^{47}$ ergs s^{-1} and $\nu L_{\nu, 300} = 4.75 \times 10^{47}$ ergs s^{-1} . Using bolometric corrections of 3.8 and 5.6 as estimated in Richards et al. (2006), we find two values for the bolometric luminosity of $L_{bol} = 1.8 \times 10^{41}$ W and 2.66×10^{41} W. We adopt the latter value as all wavelengths above 200-nm rest frame support it, it corresponds to $L_{bol}/L_{\odot} = 6.95 \times 10^{14}$ and an absolute bolometric magnitude $M_{bol} = -32.36$.

Thus, our new object would be more luminous than QSO J0100+2802 and ELIRG J2246–0526, although bolometric corrections remain uncertain. However, as Figure 4 shows, the monochromatic luminosity of ELIRG J2246–0526 exceeds that of J2157–3602 at restframe 4μ (W4-band) and most probably longwards of that. We note that the lensed QSO APM 08279+5255 at $z = 3.9$ has a magnified luminosity exceeding that of our object, but a lower intrinsic luminosity (Ibata et al. 1999).

Lacking a clearly rendered CIV line, we cannot get a reliable estimate of the black-hole mass at the moment. For the two most luminous $z > 5$ QSOs, black-hole masses were estimated from their Mg II lines: J0306+1853 at $z = 5.36$ is powered by a black hole of 10 billion solar masses (Wang et al. 2015) and J0100+2802 at $z = 6.3$ has a black hole of 12 billion solar masses (Wu et al. 2015). Both of these objects appear to radiate at the Eddington limit, and if J2157–3602 does the same, then we expect its black hole to have 20 billion solar masses.

² <http://www.astro.ucla.edu/~wright/CosmoCalc.html>

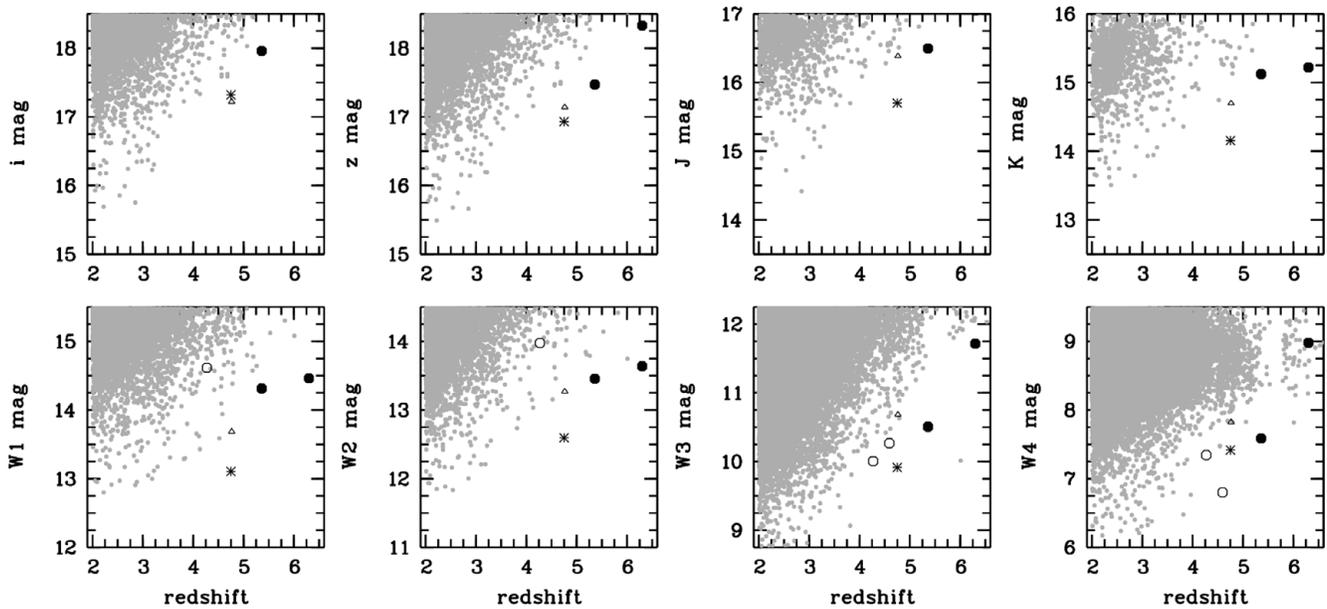


Figure 4. Photometry of J2157–3602 (asterisk) at $z = 4.75$ compared to known QSOs and ELIRGs (grey dots, from Pâris et al. 2017; Wang et al. 2016; Tsai et al. 2015); we highlight PMN J1451–1512 (triangle), the most luminous IR-bright AGN from Lacy et al. (2013) at $z = 4.27$ and from Tsai et al. (2015) at $z = 4.59$ (open symbols), and the UV-bright QSOs from Wang et al. (2015) at $z = 5.36$ and Wu et al. (2015) at $z = 6.3$ (filled symbols). Magnitudes in i - and z -bands are from SDSS and SDSS-like VST.

4 SUMMARY

ESA’s *GAIA* mission, with its unprecedented precision of proper motion measurements, has allowed us to overcome the main obstacle for the discovery of ultra-luminous QSOs in the early universe, which is the staggering contamination of candidate lists by cool stars in our own Galaxy. By combining data from *GAIA*, SkyMapper, and WISE, we identified SMSS J215728.21–360215.1 as the brightest $z \sim 5$ QSO known at this point, it has a magnitude of $z_{\text{VST}} = 16.93$ and $H_{2\text{MASS}} = 14.84$. With a redshift of $z = 4.75$, this translates into an absolute magnitude of $M_{300,\text{AB}} = -30.12$ and a bolometric luminosity of 6.95×10^{14} solar luminosities.

It is not detected in the radio surveys NVSS and SUMSS and thus not likely to be boosted in optical brightness by a relativistic jet. It appears as a single isolated point source in all images and catalogues we considered, most notably the *GAIA* DR2 database and infrared images from the VHS, and thus not likely to be strongly gravitationally lensed. The derived luminosity is likely intrinsic, making J2157–3602 the QSO with the highest unlensed UV-optical luminosity discovered so far. Our current estimate for the mass of its black hole is 20 billion solar masses, assuming it accretes at the Eddington limit.

While objects of this luminosity are exceedingly rare in the Universe, they are particularly valuable as bright background and reference sources in order to study the properties of intervening matter along the line-of-sight, and for directly probing the expansion of our Universe with new instruments in the coming decades.

Finding high-redshift QSOs is even more efficient when the follow-up of candidates is embedded in a large spectroscopic survey. Later this year, the hemispheric Taipan survey will commence at Siding Spring Observatory and take spectra of more than a million galaxies (da Cunha et al. 2017). A programme that piggybacks on Taipan aims to go after the 10^4 best candidates and will confirm a complete sample of supermassive black holes with luminous accretion in the early Universe at $4 < z < 7$.

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REFERENCES

- Abell, G. O., Corwin, H. G. Jr., & Olowin, R. P. 1989, *ApJS*, 70, 1
- Bock, D. C.-J., Large, M. I., & Sadler, E. M. 1999, *AJ*, 117, 1578
- Bromm, V., & Loeb, A. 2003, *ApJ*, 596, 34
- Childress, M. J., Vogt, F. P. A., Nielsen, J., & Sharp, R. G. 2014, *Ap&SS*, 349, 617
- Chiu, K., et al. 2005, *AJ*, 130, 13
- Condon, J. J., et al. 1998, *AJ*, 115, 1693
- da Cunha, E., et al. 2017, *PASA*, 34, e047
- Diamond-Stanic, A. M., et al. 2009, *ApJ*, 699, 782
- Dopita, M., et al. 2010, *Ap&SS*, 327, 245
- Fan, X., Carilli, C. L., & Keating, B. 2006a, *ARA&A*, 44, 415
- GAIA* Collaboration, et al. 2018, arXiv:1804.09365
- Hook, I. M., McMahon, R. G., Shaver, P. A., & Snellen, I. A. G. 2002, *A&A*, 391, 509
- Ibata, R. A., Lewis, G. F., Irwin, M. J., Lehár, J., & Totten, E. J. 1999, *AJ*, 118, 1922
- Kellermann, K. I., Sramek, R., Schmidt, M., Shaffer, D. B., & Green, R. 1989, *AJ*, 98, 1195
- Lacy, M., et al. 2013, *ApJS*, 208, 24
- Liske, J., et al. 2008, *MNRAS*, 386, 1192
- McMahon, R. G., et al. 2013, *Msngr*, 154, 35
- Oke, J. B., & Gunn, J. E. 1983, *ApJ*, 266, 713
- Pacucci, F., Volonteri, M., & Ferrara, A. 2015, *MNRAS*, 452, 1922
- Pâris, I., et al. 2017, *A&A*, 597, A79
- Qi, Z., et al. 2015, *AJ*, 150, 137
- Richards, G. T., et al. 2002, *AJ*, 123, 2945
- Richards, G. T., et al. 2006, *ApJS*, 166, 470
- Ryan-Weber, E. V., et al. 2009, *MNRAS*, 395, 1476
- Schmidt, M. 1963, *Nature*, 197, 1040
- Seymour, N., et al. 2007, *ApJS*, 171, 353
- Shanks, T., et al. 2015, *MNRAS*, 451, 4238
- Shen, Y., et al. 2011, *ApJS*, 194, 45
- Simcoe, R. A., et al. 2011, *ApJ*, 743, 21
- Skrutskie, M. F., et al. 2006, *AJ*, 131, 1163
- Tonry, J. L., et al. 2018, *PASP*, 130, 064505
- Tsai, C.-W., et al. 2015, *ApJ*, 805, 90
- Venemans, B. P., et al. 2007, *MNRAS*, 376, L76
- Wang, F., et al. 2015, *ApJ*, 807, L9
- Wang, F., et al. 2016, *ApJ*, 819, 24
- Wolf, C., et al. 2018, *PASA*, 35, 10 (doi: [10.4225/41/593620ad5b574](https://doi.org/10.4225/41/593620ad5b574))
- Wright, E. L., et al. 2010, *AJ*, 140, 1868
- Wu, X.-B., et al. 2015, *Nature*, 518, 512